

RUSSIAN ACADEMY OF SCIENCES
THE CENTRAL ASTRONOMICAL OBSERVATORY AT PULKOVO

LIVSHITS ILYA

**LARGE-SCALE MAGNETIC FIELDS
AND THE ACTIVITY OF LATE-TYPE STARS**

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The Abstract

Magnetic fields determine many physical processes in the universe. Magnetic field influences a star throughout its life. It is the reason of a complicated set of solar-type activity phenomena, which are widely spread among the low-mass stars. On the Sun, these phenomena are mainly related to the evolution of local magnetic fields. Large-scale magnetic fields on the Sun (comparable in size scale to the solar radius) determine the formation of coronal holes and streamers, and thus of solar wind (high-speed fluxes originate from coronal holes, low-speed ones are from streamers). These fields have an influence on the development of such long-duration non-stationary processes in solar corona as coronal mass ejections (CME) and flares. On the Sun, their influence is weaker than it can occur on more active stars.

The dissertation develops the idea that large-scale magnetic fields play more significant role in the formation of activity on late-type stars rather than local magnetic fields. Observations of active late-type stars with several space missions: EUVE, BeppoSAX and Chandra, are used as a basis for our study. Solar magnetic fields data obtained from ground-based observations at Wilcox Solar observatory at Stanford University.

In Chapter 1, an analysis of the total magnetic field of the Sun ('Sun-as-a-star') over several cycles of activity is performed. This analysis confirms that the rotational modulation (over years) of a similar signal from a star can serve as a good characteristic to describe the behavior of large-scale magnetic fields on active late-type stars.

In Chapter 2, a numerical MHD-simulation of long-term X-ray flares in late-type subgiants is carried out. It demonstrates that these flares exist as long as significant amount of energy arrives to the upper part of the giant magnetic loop system. It allows us to reliably determine physical conditions of plasma and size of the soft X-ray source. A new method for estimating the energy of powerful non-stationary processes on active late-type stars is proposed. It is based on determining the variation of component of the current of a large-scale magnetic field.

In Chapter 3, it is performed an assessment of the mass loss by active components of chromospherically active RS CVn-type binary systems, based on the interpretation of data on excess radiation in Fe XVIII – Fe XXIII ion lines in the extreme ultraviolet and soft X-ray and on the solving the stellar wind collision problem. The rate of quasi-stationary and non-stationary outflow in this case is 3-4 orders of magnitude greater than the solar mass loss and can reach $10^{-10} M_{\odot}$ per year.

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